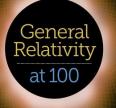
A new perspective on Einstein's cosmology

Cormac O'Raifeartaigh

Waterford Institute of Technology Dublin Institute for Advanced Studies





Brendan McCann, Michael O'Keeffe (Waterford IT) Simon Mitton (Cambridge Univ.), Werner Nahm (Dublin IAS)



Dublin Institute for Advanced Studies

75 years of DIAS

Founded in 1940 (de Valera) Modelled on Princeton Research Institute (IAS)

Two schools

School of Theoretical Physics School of Celtic Languages Only pen and paper required

Major centre for relativity

Erwin Schrödinger, J.L. Synge, K. Lanczos, L. O'Raifeartaigh Pirani, Bertotti,...









CLARENDON PRESS - OXFORD

Overview

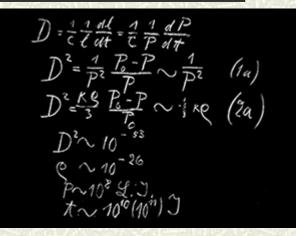
Introduction

Einstein's cosmological considerations 1915-1930

1. Einstein's cosmic model of 1931

Some new observations

- 2. Einstein's steady-state model Unfinished symphony
- **3. Einstein's cosmology review of 1933** *Further development of the Einstein-de Sitter model*
- **#** Conclusions



The Oxford blackboard (1931)



Einstein's cosmology (1915-1930)

d

The general theory of relativity (1915, 1916)

A relation between spacetime and matter Gravity = curvature of spacetime

Principles

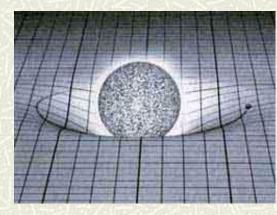
Principle of equivalence Mach's principle: no inertia relative to space Matter has a primary role, space has a derived one

- A burning question: a consistent model of the Cosmos?
 "Can relativity be followed through to the finish?"(1918)
- **Cosmological considerations (1917)** Result of long deliberation s : "A rough and winding road"

$$s^{2} = -c^{2}dt^{2} + dx^{2} + dy^{2} + dz^{2}$$
$$ds^{2} = g_{\mu\nu}dx^{\mu}dx^{\nu}$$



$$\boldsymbol{R}_{\mu\nu} - \frac{1}{2}g_{\mu\nu}\boldsymbol{R} = -\kappa \boldsymbol{T}_{\mu\nu}$$



Einstein's model of the Static Universe

Apply general relativity to the Universe (1917) Ultimate test for new theory of gravitation

Assumptions

Static universe (small velocities of the stars) Mach's principle (metric tensor to vanish at infinity) Isotropy and homogeneity (simplicity)

Boundary problem

A cosmos of closed curvature No consistent solution

Field equations modified!
 Additional term in GFE (1916)
 Radius and density defined by λ

$$R_{\mu\nu} - \frac{1}{2}g_{\mu\nu}R = -\kappa T_{\mu\nu}$$

$$\boldsymbol{R}_{\mu\nu} - \frac{1}{2}g_{\mu\nu}\boldsymbol{R} - \lambda g_{\mu\nu} = -\kappa \boldsymbol{T}_{\mu\nu} \qquad \boldsymbol{\lambda}$$

al., The Principle of Relativity (Dover, 1952), pp. 175–188. $T \text{ is well known that Poisson's equation}_{\nabla^{2}\phi = 4\pi K\rho} (1)$ in combination with the equations of motion of a material point is not as yet a perfect substitute for Newton's theory of action at a distance. There is still to be taken into account the condition that at spatial infinity the potential ϕ tends T = 0

Doc. 43

Cosmological Considerations in the General Theory of Relativity

This translation by W. Perrett and G. B. Jeffery is reprinted from H. A. Lorentz et

Some key quotes (Einstein 1917)

"In a consistent theory of relativity, there can be no inertia relative to "space", but only an inertia of masses relative to one another"

"I have not succeeding in formulating boundary conditions for spatial infinity. Nevertheless, there is still a way out...for if it were possible to regard the universe as a continuum which is finite (closed) with respect to is spatial dimensions, we should have no need at all of any such boundary conditions"

"The most important fact that we draw from experience as to the distribution of matter is that the relative velocities of the stars are very small compared with the velocity of light..... There is a system of reference relative to which matter may be looked upon as being permanently at rest "

"However, the system of equations ..allows a readily suggested extension which is compatible with the relativity postulate... For on the left hand side of the field equation...we may add the fundamental tensor $g_{\mu\nu}$, multiplied by a universal constant, $-\lambda$, at present unknown, without destroying the general covariance "

Schroedinger's comment (1918): Einstein's response (1918)

Einstein vs de Sitter

Alternative solution of the GFE *A universe empty of matter (1917)*

Solution B

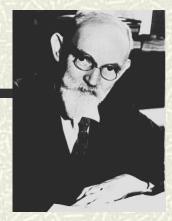
Cosmic constant proportional to curvature of space

Disliked by Einstein

Problems with singularities? (1918) Conflict with Mach's principle

The de Sitter confusion

Static or non-static - a matter of co-ordinates? Weyl , Lanczos, Lemaître



[p. 270] 5. "Critical Comment on a Solution of the Gravitational Field Equations Given by Mr. De Sitter"

[Einstein 1918c]

is

SUBMITTED 7 March 1918 PUBLISHED 21 March 1918

IN: Königlich Preußische Akademie der Wissenschaften (Berlin). Sitzungsberichte (1918): 270–272.

Herr De Sitter, to whom we owe deeply probing investigations into the field of the general theory of relativity, has recently given a solution for the equations of gravitation¹ which, in his opinion, could possibly represent the metric structure of the universe. However, it appears to me that one can raise a grave argument against the admissibility of this solution, which shall be presented in the following. The De Sitter solution of the field equations

$$G_{\mu\nu} - \lambda g_{\mu\nu} = -\kappa T_{\mu\nu} + \frac{1}{2} g_{\mu\nu} \kappa T \tag{1}$$

Prediction of redshifts – Slipher effect?

 $\lambda = 3/R$

Einstein vs Friedman

Allow time-varying solutions (1922)

Assume homogeneity, isotropy, positive curvature Two independent differential equations from GFE

Evolving universes

Density of matter varies over time

Overlooked by community

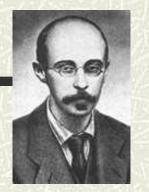
Considered 'suspicious' by Einstein $C_{J_a} \vee A^{-1}$ Mathematical correction, later retracted "To this a physical reality can hardly be ascribed"

H Negative spatial curvature (1924)

Cosmic evolution, geometry depends on matter Overlooked by community

$$\frac{3{R'}^2}{R^2} + \frac{3c^2}{R^2} - \lambda = \kappa \, c^2 \rho,$$

$$\frac{{R'}^2}{R^2} + \frac{2RR''}{R^2} + \frac{c^2}{R^2} - \lambda = 0.$$
 Alexander Friedman (1888 - 1925)



$$\frac{1}{c^2} \left(\frac{\mathrm{d}R}{\mathrm{d}t}\right)^2 = \frac{A - R + \frac{\lambda}{3c^2}R^3}{R}$$

$$t = \frac{1}{c} \int_{a}^{R} \sqrt{\frac{x}{A - x + \frac{\lambda}{3c^{2}}x}} \, \mathrm{d}x + B$$

Einstein vs Lemaître (1927)

Expanding model of the Universe from GR
 Similar but not identical to Friedman 1922
 Starts from static Einstein Universe

$$2\frac{R''}{R} + \frac{R'^2}{R^2} + \frac{1}{R^2} = \lambda - \kappa p$$

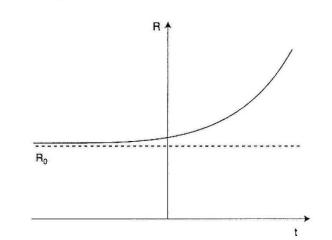
 $3\frac{\mathbf{R}^{\prime 2}}{\mathbf{R}^2} + \frac{3}{\mathbf{R}^2} = \lambda + \kappa \rho$

Fr Georges Lemaître

- Redshifts of galaxies = expansion of metric? Redshifts from Slipher, distances from Hubble H = 585 kms⁻¹Mpc⁻¹
- **#** Ignored by community

Belgian journal (in French) Rejected by Einstein: "Votre physique est abominable"

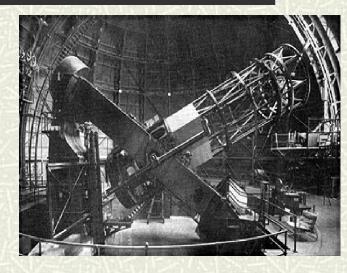
Lemaître's recollection (1958) "Einstein not up-to-date with astronomy"

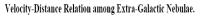


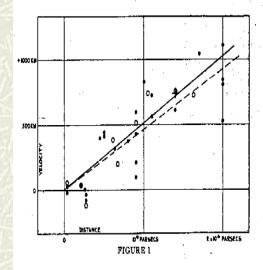


Hubble's law (1929)

- A redshift/distance relation for the nebulae?
 Motivation: establishing distances of all nebulae
- Linear relation (Hubble, 1929)
 20 redshifts from Slipher: not acknowledged
 Most important data point not shown (8Mpc, 4000 km/s)
- Landmark result in astronomy H = 500 kms⁻¹ Mpc⁻¹
- Not the expanding universe!
 Astronomy, not cosmology







The watershed

• **RAS meeting (1930)**

If redshifts are velocities, and if effect is non-local Hubble's law = expansion of space? (Edd., de Sitter)

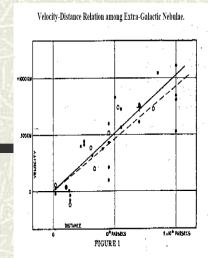
Dynamic model required

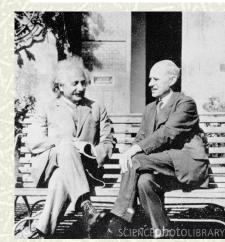
"Motion in Einstein's model or matter in de Sitter's?"

Lemaître's intervention 1927 expanding model republished in English (1931) Observational section omitted (rightly)

Lemaître model circulated

Time-varying radius, density of matter Friedman 1922 models become known Positive curvature





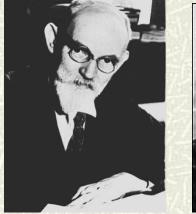


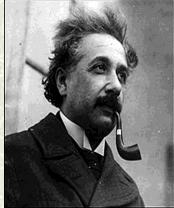
Models of the evolving universe (1930 -)

- **Tolman (1930, 31)** *Expansion caused by annihilation of matter ?*
- Eddington (1930, 31)
 On the instability of the Einstein universe Expansion caused by condensation?
- de Sitter (1930, 31)
 Variety of expanding models
- Heckmann (1931,32)
 Spatial curvature (not translated)
- Einstein (1931, 32) JMK Friedman-Einstein model $\lambda = 0, k = 1$ Einstein-de Sitter model $\lambda = 0, k = 0$









<u>If</u> redshifts represent expansion... I<u>f</u> effect is global...

Energies Zum boundagierten Problem der ellgenetiem Behrbeitendesein 233

Zum kosmologischen Problem der allgemeinen Relativitätstheorie.

Unite dana hoomologiethen Problem with the Page these the Boeringfragher des Binanas in agodien und diver dio Art des Verrichang der Mosteis im gesten verstaation, webeit die Materie des Stears and Stenregetung auf gesten verstaation, webeit die Materie des Stears auf die entgewennen die Stears auf die Stears auf die Stears auf die entgewennen beher übereit dieses Problem ist Ausgefür nation, dass die entgewennen die Housen Arsteinung die Gestenre auf die entgewennen der Housenstkerten über dieses Frechtung und verschlauser, aus der einstein Anzeitsen Biefer Ubsiehen nach zum einschlauser, ausdaher ist ein Housen Anzeitsen Biefer Ubsiehen nach gestenret, auch die Versteilung der eventrage dicharten Mehr Ubsiehen net beite gestenret, nicht der Thusteis aum Worg inflations Mehr Ubsiehen net beite gestenret, nicht der Thusteis aum Worg

In meiner urspränglichen Untersochung ging ich von folgenden Anstrasen nus:

 Alto Stellen des Universituss sind gleichwertig; im speziellen mit aleo auch die ürtlich gemittalte Dichte der Stermunierie überalt gleich sein.

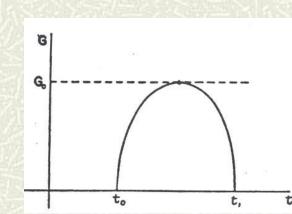
 Municheke Structur und Dieldie solllen zwillen konstruit sein. Iheinik zeigen ich, dall nam beiden Annahmen und einer von Null verschiedenen m
ürtigene Glied an die Feligbeiehungen der allgemeinen Relativit
ütstesorie

 $(R_n - \pm y_n R) \pm \pm y_n = -\pm T_n \dots$

issue tilhichungen wird durch eine ninndheit sphirische sonische Weit vom sdus $P = |V|^{\frac{3}{2}}$ Genüge gelrister, wenn z die (drachfreis) mittiere Dichte der

Interio Industri. Neukolem um shere dwerb Brewnas Rennhaus kier gewonden ist, shali die obersphältstehen. Nehet gleichnofötig über den Raum vertreilt and in einer hantstruchversegenge perpförs sich dywingbreis selerer um dezen gestematische obverseiheltetingen als Degelterfichko en donten hiet), het die Anathme (r) uter vitikische Nitzer des Buzues leiten Bertaufung anhlin, mai des entchen die Franze, ob die utgenotien Keintreithreihenzie von diesem Befunden orderaschlöt on geben verenze.

 $(\frac{dP}{dt})^2 = c^2 \frac{P_0 - P}{P}$



Einstein's 1931 model (F-E)

- Einstein's 'first' model of the Expanding Universe Occasionally cited, rarely read (not translated)
- # Adopts Friedman 1922 model

Time-varying, closed universe: k = 1

Set cosmic constant to zero
 Instability of static solution
 Hubble's observations

$$D = \frac{1}{P} \frac{dP}{dt} \cdot \frac{1}{c} \qquad D^2 =$$

Extraction of cosmic parameters! $P \sim 10^8$ lyr : $\rho \sim 10^{-26}$ g/cm³ $t \sim 10^{10}$ yr : conflict with astrophysics Attributed to simplifying assumptions (homogeneity)

 $\frac{3P'^2}{P^2} + \frac{3c^2}{P^2} - \lambda = \kappa c^2 \rho .$ $\frac{P'^2}{P^2} + \frac{2P''}{P} + \frac{c^2}{P^2} - \lambda = 0$

$$\frac{1}{c^2} \left(\frac{\mathrm{d}R}{\mathrm{d}t}\right)^2 = \frac{A - R + \frac{\lambda}{3c^2}R^3}{R}$$

$$= \frac{1}{P^2} \frac{P_0 - P}{P} \qquad P \sim \frac{1}{D}$$
$$D^2 = \frac{1}{3} \kappa \rho \frac{P_0 - P}{P}$$

 $D^2 \sim \kappa \rho$

4 D

Von A. EINSTEIN

Einstein's 1931 model revisited

First translation into English O'Raifeartaigh and McCann 2014

$$D = \frac{1}{P} \frac{dP}{dt} \cdot \frac{1}{c}$$

$$D^2 = \frac{1}{P^2} \frac{P_0 - P}{P}$$

Anomalies in calculations of radius and density *Einstein:* $P \sim 10^8$ lyr, $\rho \sim 10^{-26}$ g/cm³, $t \sim 10^{10}$ yr *We get:* $P \sim 10^9$ lyr, $\rho \sim 10^{-28}$ g/cm³, $t \sim 10^9$ yr

Source of error?

Oxford blackboard: $D^2 \sim 10^{-53} \text{ cm}^{-2}$ should be 10^{-55} cm^{-2} Time miscalculation $t \sim 10^{10} \text{ yr}$ (should be 10^9 yr) Non-trivial error: misses conflict with radioactivity

Not a cyclic model

"Model fails at P = 0" Contrary to what is usually stated $D^2 = \frac{1}{3}\kappa\rho \frac{P_0 - P}{P} \qquad D^2 \sim \kappa\rho$

 $P \sim \frac{1}{R}$

Oxford lecture (May 1931)

 $D = \frac{1}{C} \frac{1}{C} \frac{dl}{dt} = \frac{1}{C} \frac{1}{P} \frac{dl}{dt}$ $D^2 = \frac{1}{P^2} \frac{P_c - P}{P} \sim$ $\underline{P_{s'-P}} \sim$

Some key quotes (Einstein 1931)

"The cosmological problem is understood to concern the question of the nature of space and the manner of the distribution of matter on a large scale, where the material of the stars and stellar systems is assumed for simplicity to be replaced by a continuous distribution of matter."

"Now that it has become clear from Hubbel's results that the extra-galactic nebulae are uniformly distributed throughout space and are in dilatory motion (at least if their systematic redshifts are to be interpreted as Doppler effects), assumption (2) concerning the static nature of space has no longer any justification...."

"Several investigators have attempted to account for the new facts by means of a spherical space whose radius *P* is variable over time. The first to try this approach, uninfluenced by observations, was A. Friedman,¹ on whose calculations I base the following remarks."

"However, the greatest difficulty with the whole approach... is that according to (2 a), the elapsed time since P = 0 comes out at only about 10^{10} years. One can seek to escape this difficulty by noting that the inhomogeneity of the distribution of stellar material makes our approximate treatment illusory."

Einstein Archives Online Albert Einstein

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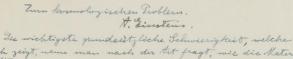
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sich zeigt, nem man nach der tit fragt, wie die Materie then Raum in schr grossen Dimensionen erfieltt, liegt bekanntlich durin, dues die Granitationsgesethe im telgemeinen mit der Rypothese einer endlichen mittleren Fichte der Materie micht verträglich sind. Ichon zu der Zeit, als man woch allymein an Newtons gravitations - Theorie furtheelt, hat deshalt Teeliger das Newton'selve Gesetz durch time tostandstruchtion modifiziert, welche fits grosse Hostinde & erheblich schweller abfielt als 72

Auch in der allgemeinen Relativitätstheorie titt diese Tehnonigkest unt. Ich habe aber früher gezeigt, dass letytere durch Sinf showing des sogenannten, & - Gloedes

Einstein Archives Online

Alberg Timeteins



English V

Home Zum kosmologischen Problem. Related Items

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Zum kosmologischen Problem.

Archival Call Number:	2-112
Begin Date:	1931-01-01
End Date:	1931-12-31
Main Author:	Einstein, Albert (Author)
Language:	German
Archival Location:	Albert Einstein Archives, The Hebrew University of Jerusalem, Israel
Number of Pages:	4



Items related to Readex number 205 of the bibliographical list of Einstein's Writing

- 1. vor allem mein Compliment zu Ihrem Büchlein (13-308)
- 2. Vor allem mein Kompliment zu Ihrem Buechlein (13-309)
- Zum kosmologischen Problem der Allgemeinen Relativitätstheorie. (4-53)

Kurze Zusammenfassung. (4-54)

Die Gleichungen (1) liefor - 2 a2 + de2 = 0 3 x2 - 1c2=xec2

rder

Die Victite ist also konstant und bestimment die Sepansion tes auf das Vorgeichen.

"The density is thus constant and determines the expansion"

 $\alpha^2 = \frac{\kappa_c^2}{\kappa_c^2} g \cdots (4)$

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Zum kosmologischen Problem der allgemeinen Relativitätstheorie. By: Einstein, Albert (Author) Archival Call Number: (81-663)

Beiträge zum

2. Einstein's steady-state model

Unpublished manuscript

Archived as draft of F-E model (1931) Similar title, opening to F-E model

Something different

Cosmological constant "The density is thus constant and determines the expansion"

Steady-state model of the Expanding Universe

Anticipates Hoyle solution Written in early 1931 Fatal flaw: abandoned

Turn kosmologischen Troblers. H. Ginstein

The wichtigete grundeetzliche Gelwerigkeit, welche sich zeigt, nem man nach der tit fragt, we die Materse aus Theigen uns schr grossen Dimensionen affelt, liegt bekanntlich darin, duss die Grantertionsgesetze im telgemeinen mit des Hypothese einer endlichen mitteren Tichte die Materie mit des Hypothese einer endlichen mitteren Tichte die Materie aust des Hypothese einer endlichen mitteren Tichte die Materie aust des Hypothese einer Grantertions-Theorie feitheelt, het allgemein an Nentono gravitations-Theorie feitheelt, het dishalt Geeliger das Neuton'sche Gesetz derektente fürstendstuckten modefiziert, welche für geosse Absteinde reicheltete scheelte alföllt als z.

Auch in der allgemeinen Relativitätetheorie teitt diere Gehvörigkeit auf, Ich habe aber friher gezigt, dass litztere durch Imfährung des sogenennten, d-Geledes" in die Feldgleichungen überminden werden kann. Fie Feldgleichungen können dann in der Toren geselwieben werden

 $(\mathcal{R}_{ik} - \frac{1}{2}g_{ik}\mathcal{R}) - dg_{ik} = \kappa \mathcal{T}_{ik} \cdots (l)$

$$\frac{\Im}{2} = \frac{3}{2} \alpha^{2} + \lambda c^{2} = 0$$

$$\frac{\Im}{2} \alpha^{2} + \lambda c^{2} = 0$$

$$\frac{\Im}{2} \alpha^{2} + \lambda c^{2} = 0$$

$$\frac{\Im}{2} \alpha^{2} - \lambda c^{2} = \kappa \rho c^{2}$$

$$3\alpha^{2} / 4 - \lambda c^{2} = \kappa \rho c^{2}$$

$$\alpha^{2} = \frac{\kappa c^{2}}{3} \rho c^{2} + \kappa c^{2} \rho c^{2}$$

$$\alpha^{2} = \frac{\kappa c^{2}}{3} \rho^{2} + \kappa c^{2} \rho^{2}$$

oder

 $\alpha' = \frac{\kappa c^2}{3} \varphi \cdots (4) \qquad \alpha^2 = \frac{\kappa c^2}{3} \rho$ die Gielete ist also konstant und bestimmt die Zepansion des auf das Vorgeichen.

Einstein's steady-state model (Jan 31)

Problem with evolving models

"De Sitter and Tolman have already shown that there are solutions to equations (1) that can account for these [Hubbel's] observations. However the difficulty arose that the theory unvaryingly led to a beginning in time about $10^{10-} 10^{11}$ years ago, which for various reasons seemed unacceptable."

New solution

"In what follows, I wish to draw attention to a solution to equation (1) that can account for Hubbel's facts, and in which the density is constant over time..

If one considers a physically bounded volume, particles of matter will be continually leaving it. For the density to remain constant, new particles of matter must be continually formed within that volume from space "

Mechanism

"The conservation law is preserved in that, by setting the λ -term, space itself is not empty of energy; its validity is well known to be guaranteed by equations (1)."

An abandoned model

Correct geometry

de Sitter metric

- **#** Simultaneous equations
 Eliminate λ
 Relation between α² and ρ
- Einstein's crossroads
 Null solution on revision
 Tolman? (Nussbaumer 2014)
 Declined to amend GFE
- **#** Evolving models

Less contrived: set $\lambda = 0$

Fur Nachfolgenden will sich auf eine Lösung der Gleichung (1) aufmächstem machen, welche Hubbel's Thatsachen gerecht wird, und in welcher die Dichte pettich konstant ist. Dere Lösung ist zwar in dem allgemeinen Schema Tolman's withalten, rehesst aber hisher wicht in Betracht gezogen worden zu seen. 1 Jeh setze au

$$ds^{2} = -e^{at} (dx_{1}^{2} + dx_{2}^{2} + dx_{3}^{2}) + c^{2} dt^{2} \cdots dt^{2})$$

oder

Die Dichte ist also konstant und bestimmet die Expansion Les auf das Vorgeüchen.

Der behaltnurgesatz bleebt deedurch zuwahrt, dess bei Tetpung des 1-Gludes die Kamme selbst nicht energetisch leer ist; sime Gelting wird bekanntlich durch die Gletchungen (1) gewährleistet.

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NATURE | NEWS

Einstein's lost theory uncovered

Physicist explored the idea of a steady-state Universe in 1931.

Davide Castelvecchi

24 February 2014

New Discovery Reveals Einsteir Tried To Devise A Steady State Model Of The Universe

2 comments, 2 called-out + Comment Nov + Follow Comments

Almost 20 years before the late Fred Hoyle and his colleagues devised the Steady State Theory, Albert Einstein toyed with a similar idea: that the universe was eternal, expanding outward with a consistent input of spontaneously generating matter.

An Irish physicist came across the paper last year and could hardly believe According to this week's article in Nature,

model of the universe very different to today's Big Bang Theory.

The manuscript, which hadn't been referred to by scientists for decades.

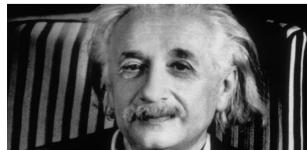




Feb 25, 2014 | By Davide Castelvecchi and Nature magazine



Scientists uncovered misfiled papers while searching Jerusalem university's online archive



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- 08:42 Gardaí investigate death of woman in Dublin
- 08:25 Flannery faces call from all parties to attend PAC

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A useful find

H New perspective on steady-state theory (1950s)

Logical idea: not a crank theory Tolman, Schroedinger, Mimura : considered steady-state universe

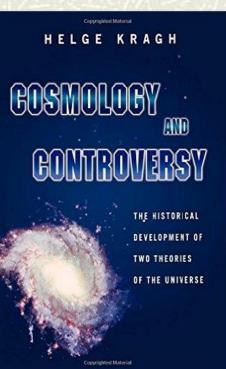
Insight into scientific progress

Unuccessful theories important in the development of science

Links with modern cosmology Creation energy and λ: dark energy de Sitter metric: cosmic inflation

Insight into Einstein's cosmology

Turns to evolving models rather than introduce new term to GFE Pragmatic approach: F-E model



3. Einstein-de Sitter model (1932)

Curvature not a given in dynamic models (Heckmann)

Not observed empirically Remove spatial curvature (Occam's razor)

 $ds^2 = -R^2(dx^2 + dy^2 + dz^2) + c^2dt^2$

Simplest Friedman model

Time-varying universe with $\lambda = 0$, k = 0, p = 0Estimate of density : $\rho = 10^{-28} \text{ g/cm}^3$ Important hypothetical cosmos: critical case

Becomes standard model

Despite high density of matter, age problem Time evolution not considered Eddington's anecdote

$$\frac{1}{R^2} \left(\frac{dR}{cdt}\right)^2 = \frac{1}{3} \kappa \rho.$$

 $h^2 = \frac{1}{2} \kappa \rho$



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Volume 1	8
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Number 3

ON THE RELATION BETWEEN THE EXPANSION AND THE MEAN DENSITY OF THE UNIVERSE

March 15, 1932

By A. EINSTEIN AND W. DE SITTER

Communicated by the Mount Wilson Observatory, January 25, 1932

In a recent note in the *Göttinger Nachrichten*, Dr. O. Heckmann has pointed out that the non-static solutions of the field equations of the general theory of relativity with constant density do not necessarily imply a positive curvature of three-dimensional space, but that this curvature may also be negative or zero.

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Einstein-de Sitter model revisited

Einstein's cosmology review of 1933

Overview of dynamic models from first principles Significant exposition Culminates in Einstein-de Sitter model Including time evolution

Published in 1933

French book (transl. Solovine); small print run Intended for scientific journal; not submitted

Parameters extracted

Critical density of 10⁻²⁸ g/cm³ : reasonable Timespan of 10¹⁰ years: <u>incorrect estimate</u> Conflict with astrophysics: attributed to simplifications

$$2A \frac{d^{2}A}{dt^{2}} + \left(\frac{dA}{dt}\right)^{2} = 0$$
$$3 \left(\frac{dA}{dt}}{A}\right)^{2} = \varkappa \rho c^{2}.$$

$$3h^2 = \varkappa \rho c^2 \ (= 8\pi \mathrm{K} \rho)$$

$$\mathbf{A} = c \left(t - t_0 \right)^{\frac{2}{3}}.$$

und Last de m-relateristischen Physik abolist" neurons, or hat does folgende Bedentung, Turstens hat dort A Tait (in demeetber Jun die Bedentung von sine Realitet on die Masse. Die Hoordensten enbeging unf das gemeihlte Begings ustern bedenters muni Gelbar Messegebusser. Totze des Geometrie und Kinematik bedeuten deshalt Telationen zwischen Messengen, welche die Bedentung von physikalischen Behauptungen huben, die nicht. ader falech sein hömmen. Das Instialsystem federtet ame Realität. mail sins Wahl in das Trogheitegesetz singeht. Tweitens ist dies Hycikalisch Reule, may mit den Korten Raum , Test bezeichnet wird gesetzmanigkeiten unabhängig om dem Tuhalter des ibrigen physikalesch - Reales d. h. Timethängig van den Herper-Do Tubegriff die Bezichungen griescher Merreraltaten, die allein au Marstofue and Whales you germine wind, ist mark deen There is me die Testiling und Beneging die Herper mathangig, ebenes das Tuntialization. For physis' Rame ist gentermanne physichall is inkend abor wicht physikalisch beeinflussbar

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SUR LA STRUCTURE COSMOLOGIQUE DE L'ESPACE ⁽¹⁾

Si nous appelous l'espace et le temps de la physique prérelativiste « absolus », il faut y voir la signification suivante. Tout d'abord l'espace et le temps et, par suite, le système de référence, y figurent dans le même sens comme réalité que, par exemple, la masse. Les coordonnées du système de référence choisi y correspondent immédiatement à des résultats de mesure (?). Les propositions de géométrie et de cinématique signifient pour cette raison des relations entre des mesures avant la valeur d'àfitmations physiques, qui peuvent être vraies ou fausses. Le système d'inertie possède une réalité physique, parce que son choix entre dans la loi d'intrite. En second leu, cette réalité physique, qui et désignée par les termes espace + temps, est, quant à ses lois, indépendante du comportement des autres réalités physiques, par exemple, des corps.

Some key quotes (Einstein 1933)

"Since, according to the general theory of relativity, the metric properties of space are not given in themselves but are instead determined by material objects that force a non-Euclidean character on the continuum, a problem arises....since we may assume that the stars are distributed with a finite density everywhere in the world, that is, a non-zero average density of matter in general, there arises the question of the influence of this mean density on the (metric) structure of space on a large scale; this is the so-called cosmological problem "

"Thus the theory can now, without the introduction of a λ -term, accommodate a finite (mean) density of matter ρ on the basis of equations (1), using relation 3(a) with *P* (and ρ) variable over time"

"It follows from these considerations that in the light of our present knowledge, the fact of a non-zero density of matter need not be reconciled with a curvature of space, but instead with an expansion of space. Of course, this does not mean that such a curvature (positive or negative) does not exist. However, there is at present no indication of its existence. In any case, it may well be substantially smaller than might have been suggested by the original theory (see equation 5)."

Einstein's cosmology: conclusions

Major test for general relativity

Conscious of assumptions of homogeneity, isotropy

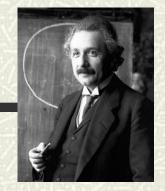
Embraces dynamic cosmology

New evidence – new models (JMK) Timespan of expanding models puzzling Steady-state universe?

Evolving models (less contrived)

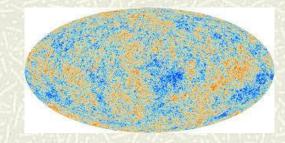
Cosmic constant not necessary Extraction of parameters compatible with observation Closed and open models Timespan problem attributed to simplifying assumptions

Verdict (1933, 1945): more observational data needed





Hubble constant revised



Cosmic microwave background Homogeneous, flat universe

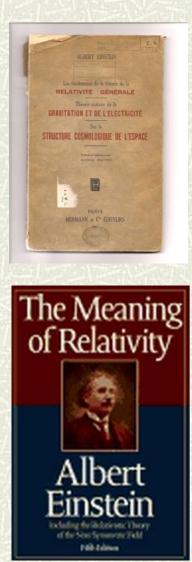
No mention of origins



Einstein's greatest hits (cosmology)

- **#** Einstein's model of the Static Universe (1917) First relativistic model of the cosmos
- **#** Einstein's steady-state model (Jan 31) Natural successor to static model: abandoned
- **Friedman-Einstein model of the Universe (1931)** Use of Hubble constant to extract observational parameters
- **#** Einstein-de Sitter model of the Universe (1932)
- **# 1933 review: 1945 review (Appendix)** *Conversations with Gamow, Godel, Straus*

No mention of origins



Further reading



Historical Perspectives on Contemporary Physics

An image of the blackboard used in Einstein's 2nd Rhodes lecture at Oxford in April 1931 (reproduced by permission of the Museum of the History of Science, University of Oxford)

Einstein's cosmic model of 1931 revisited: An analysis and translation of a forgotten model of the universe

by Cormac O'Raifeartaigh and Brendan McCann

ecpsciences



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Eur. Phys. J. H DOI: 10.1140/epjh/e2014-50011-x

THE EURO PHYSICAL

The European Physical Journal H

Einstein's steady-state theory: an abandoned model of the cosmos

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Abtract. We present a transition and analysis of an unpublished manneriph by Absent Existem in within the attempted to construct a vicasity-static model of the mirrore. The manneriph tailing appears to have been written in early 1901, demonstrates that Existen core considered a songle model is which the mean density of matter the an expanding mirrore is maintained constant by the continuous formation of matter time empty space. This model sever different to previously





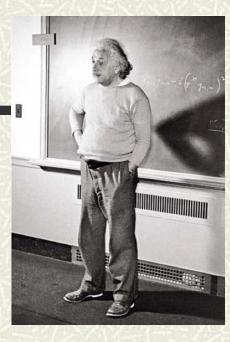
Coda: Einstein vs Hoyle

Fred Hoyle in Princeton (1952, 53)

Einstein remark to Manfred Clynes "Romantic speculation" (Michelmore 1962)

Letter to Jean Jacques Fehr (1952)

"The cosmological speculations of Mr Hoyle, which presume a formation of atoms from space, are in my view much too poorly grounded to be taken seriously. On the whole, it seems to me more reasonable to seek a solution to problems far closer to hand, e.g., the theory of quantum phenomena or the further development of the general theory of relativity. The popular literature on the subject is not very fruitful, as it encourages flights of fancy rather than clear thinking. In my opinion, this is less because of the nature of the problem itself than because our theoretical insight is still extremely deficient."



Herrn Jean Jaques Fehr c.o.Grace Y Cia Sucursal Busaaventura Colombia S.A.

Schr geehrter Herr Fchr

Die kosmologischen Spekulationen von Herrn Hoyle, welche eine Entstehung von Atomen aus den Ruum vorwussetzen, sind mich meiner Annicht viel zu wenig begründet, um ernnt genommen zu werden. Heberhaupt scheint en in vorminftiger, sich um eine Lönung weit nüherliegender Frobleme zu benühen, sich die Theorie der Quantenphennenen und der Fortentisder allgreichen Relativisitisteberie. Die populäre Literatur führ den Gegenstund ist freilich nicht schunuhrhaft, weil sie nehr die Flantenis als dan klitz Denken anzegt. Dies liegt mich meiner Annicht weniger an der Xuur der Frobleme solltet als daran, dass um ere theoretische Einsicht noch Humserst mangelhaft ist. Freumlich grüßes Sie

A. Cincline

Further reading

Einstein's 1931 model

Einstein's cosmic model of 1931 revisited; an analysis and translation of a forgotten model of the universe. O'Raifeartaigh, C. and B. McCann. 2014 *Eur. Phys. J (H)* 39(1):63-85

Einstein's steady-state manuscript

Einstein's steady-state theory: an abandoned model of the cosmos. O'Raifeartaigh, C., B. McCann, W. Nahm and S. Mitton. 2014 *Eur. Phys. J (H)* 39(3):353-367

Einstein-de Sitter model

Einstein's cosmology review of 1933: a new perspective on the Einstein-de Sitter model of the cosmos. O'Raifeartaigh, C., M.O'Keeffe, W. Nahm and S. Mitton. 2015. Eur. Phys. J (H) 40 (3) 301-335

Review paper



Observational parameters needed (1930s)

- **#** Spatial curvature k = -1, 0, 1?
- **#** Cosmic constant $\lambda = 0$?
- **# Deacceleration** $q_0 = \ddot{R}/\dot{R}^2$
- **#** Density of matter

Hubble constant

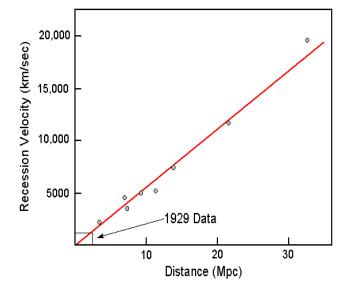
其

- $\rho < \rho_{crit}$?
- **#** Timespan $\tau = 10$

 $\tau = 10^{10} \text{ yr}?$

 $\dot{R}/R = 500 \ km s^{-1} Mp c^{-1}?$

Hubble & Humason (1931)



What do redshifts represent? Is expansion a local effect?

Hubble and Tolman 1935

Einstein's steady-state model and cosmology today

Dark energy (1998)

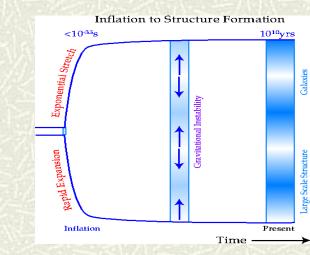
Accelerated expansion (observation) Positive cosmological constant

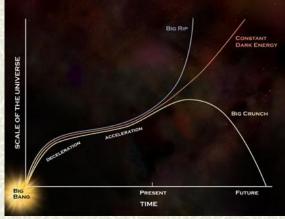
Einstein's dark energy

"The conservation law is preserved in that, by setting the λ -term, space itself is not empty of energy; its validity is well known to be guaranteed by equations (1)."

Cosmic inflation

Inflationary models use de Sitter metric Used in all steady-state models Flat curvature, constant rate of matter creation Different time-frame!





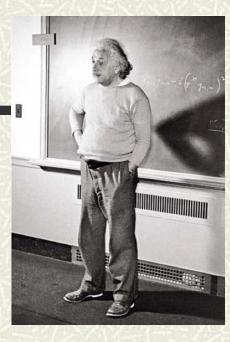
Einstein vs Hoyle

Hoyle in Princeton (1952, 53)

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Steady-state cosmology today

Observable universe not in a steady state

Evolution of galaxies Cosmic microwave background

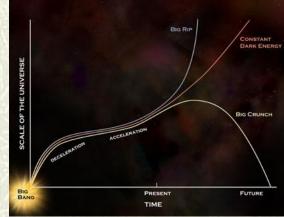
Inflationary cosmology = steady-state model

de Sitter metric Steady-state model with different time-frame! (Hoyle 1990) Matter creation term not mandatory in Hoyle models (McCrea 1951)

Eternal inflation

Different regions undergo different inflation? Inflation begets further inflation (Vilenkin 1983; Linde 1986) Observable universe embedded in global steady-state cosmos?

Hoyle's revenge! (Hoyle and Narlikar 1966; Barrow 2005)



The Scientific Legacy of **Fred Hovle**

Edited by Douglas Gough

Einstein's steady-state model and cosmology today

Accelerated expansion (1998)

Supernova measurements Dark energy – positive cosmological constant

Einstein's dark energy

"The conservation law is preserved in that, by setting the λ -term, space itself is not empty of energy; its validity is well known to be guaranteed by equations (1)." **Anticipates positive cosmological constant**

De Sitter line element

 $ds^{2} = -e^{\alpha t} \left(dx_{1}^{2} + dx_{2}^{2} + dx_{3}^{2} \right) + c^{2} dt^{2} \dots$ Necessary for all steady-state models

Identical to inflationary models (different time-frame)

Einstein's cosmology (traditional view)

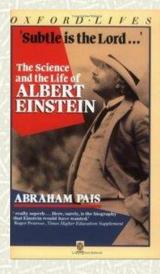
Static, bounded model of the cosmos (1917) Resistance to empty de Sitter 'static' solution

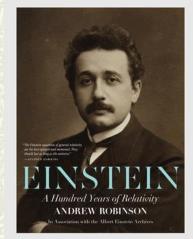
Resistance to time-varying models *Friedman, Lemaitre*

Einstein the conservative
 Hidebound by philosophical prejudice
 Disinterested in evolving cosmologies

Dynamic model 1932 (1931)
Minimalist, cursory models

A new perspective? Guided by physical intuition





Taking $T_{44} = \rho c^2$ (all other components zero) in the *time* component of equation (1) we obtain $\left(R_{44} - \frac{1}{2}g_{44}R\right) - \lambda g_{44} = \kappa \rho c^2$. This gives on analysis - $3\alpha^2 / 4 + 3\alpha^2 / 2 - \lambda c^2 = \kappa \rho c^2$ the second of Einstein's simultaneous equations.

From the *spatial* component of equation (1), we obtain $\left(R_{ii} - \frac{1}{2}g_{ii}R\right) - \lambda g_{ii} = 0$. This gives on analysis $3\alpha^2/4 - 3\alpha^2/2 + \lambda c^2 = 0$ for the first of the simultaneous equations.

It is plausible that Einstein made a sign error here, initially getting $3\alpha^2/4$ + $3\alpha^2/2 + \lambda c^2 = 0$ for this equation.